

Exotic MOTSs in Spherically Symmetric Black Holes



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based on [1] with Robie Hennigar, Liam Newhook, and Ivan Booth.

Introduction

The behaviour of the horizons during a black hole merger had not been understood. Recently, there been computational work that show how the two apparent horizons become one.



 $D \rightarrow 4$ *Gauss-Bonnet :* A Horndeski-type scalar-tensor theory obtained using dimensional reduction methods [6]:

$$S = \int d^4x \sqrt{-g} \left[R + \alpha (\phi \mathcal{G} + 4G^{\beta\gamma} \partial_\beta \phi \partial_\gamma \phi - 4(\partial \phi)^2 \phi + 2(\partial \phi)^4) \right]$$

$$f(r) = 1 + \frac{r^2}{2\alpha} \left(1 - \sqrt{1 + \frac{8\alpha M}{r^3}} \right) .$$

Left: "Pair of Pants" diagram [2] depicting the event and apparent horizons throughout a binary black hole merger.

Right: Top-down plots of numerically obtained apparent horizons at time-slices [3].

The data suggested the merger process involves selfintersecting *marginally outer-trapped surfaces* (MOTSs, summary in [4]). It was then shown that (seemingly infinite) self-intersecting MOTSs also exist in analytical black hole spacetimes, such as the Schwarzschild spacetime [5].



Plots of the 6 outer-most MOTSs in the Schwarzschild BH [6]. *Hypothesis:* These self-intersecting MOTSs will appear in similarly simple black hole spacetimes.

Set-up

 $\kappa =$





Plot of f(r) for varying α (blue). The red curve is Schwarzschild.

Results

The existence of an inner (spherical) horizon seems to put an upper bound on number of axisymmetric MOTSs.



These results are a first step to understanding evolution in black hole merger processes. Further published work have shown comparable results in spacetimes with rotation [7] and non-time-symmetric coordinates [8]. Current work investigates the behaviour of MOTSs in asymptotically AdS spacetimes.

-0.5

For an induced metric q_{AB} and null normal vector l^{α} of a surface, surface is a MOTS if it has vanishing outward expansion (by definition)

 $\Theta_{(+)} = q^{AB} e^{\alpha}_A e^{\beta}_B \nabla_{\alpha} l_{\beta} \; .$



 $\Theta_{(\ell)}=0$

Figure picturing expansion via light spheres, the null normal vector l^{α} is in red and the surface is blue.

Consider spherically symmetric black hole spacetimes, write a Painlevé-Gullstrand-type metric

 $ds^2 = -f(r) dt^2 + 2\sqrt{1 - p(r)f(r)} dt dr + p(r) dr^2 + r^2 d\Omega^2$. Only looking at axisymmetric MOTSs, q_{AB} is the orbit space and the problem boils down to solving a pair of coupled ODEs: $p/\dot{r}^2 - 2p\dot{\rho}^2 = r\dot{\rho}$

$$\begin{split} \ddot{r}(s) &= -\frac{p'\dot{r}^2 - 2r\theta^2}{2p} + \frac{r\theta\kappa}{\sqrt{p}} \ ,\\ \ddot{\theta}(s) &= -\frac{2\dot{r}\dot{\theta} + \sqrt{p}\dot{r}\kappa}{r} \ .\\ -\frac{p\dot{r}\cot(\theta) - r\dot{\theta}}{r\sqrt{p}} + \frac{rp^2\dot{r}^2f' + r\dot{r}^2p' - 2(r^2\dot{\theta}^2 + 1)(1 - pf)}{2r\sqrt{p(1 - pf)}} \end{split}$$

References

-0.2 -0.1

0.0

0.1 0.2

-0.4

-0.2

0.0

0.2

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